A SUMMARY OF THE MEETING
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I doubt that a detailed summary of this extensive meeting would
be very useful, particularly since the details appear in these proceed-
ings, so I won't attempt one. I can, however, try to put all this into
some kind of perspective, and itemize those things I think are im-
portant to our collaboration and to the science we try to do. This
is a personal view: I'm sure others would have selected differently.
The items below are not in order of importance, and only time will
tell us if any of them are important at all.

- While the high temporal resolution we can obtain with the WET
  proved to be essential for resolving, and therefore understanding,
  the patterns present in the multi-mode white dwarf pulsators, it
  is not the only instrument we can bring to bear on our targets,
  and sometimes not even the most effective one. Chris Clemens
  has made an eloquent plea for single-site data on the DAV pul-
sators, and Scot Kleinman has shown how powerful single-site
data can be in unravelling the DAV pulsators if there is enough
of it. The DAV pulsators do not reveal all of their pulsation
structure in any one WET run, nor even in one observing sea-
son. Since they are evolving more slowly than their hotter sib-
lings, the timescale on which they show differing parts of their
underlying structure is longer, so it takes many years of accu-
mulated data to see it all. So far as we can tell, though, the
underlying structure is stable and discernable: we just can't see
it all at once.

- The measurement of asteroseismological distance (via extrac-
tion of the absolute luminosity by model fitting) seems to work,
and work very well. The distances to PG 1159, GD 358 and
G 29-38 agree with the distance derived from parallax measure-
ments, and so far there have been no failures of the method. It
was given additional credibility at this meeting by the very nice
agreement on the distance to PG 2131: the asteroseismologi-
cal distance derived by Kawaler, O'Brien and their colleagues,
and the independent distance derived by Howard Bond based
on main-sequence fitting to PG 2131's optical companion, a
red dwarf cleanly resolved by the Hubble Space Telescope im-
ages. Incidentally this agreement also shows that PG 2131 and
its companion are a gravitationally- bound pair – the standard
method for proving such an association requires that orbital motion be observed, and takes a long time to demonstrate. So far as I know, this is the first time such an association has been shown by two independent distance measurements: we usually consider ourselves very lucky just to get one of them.

- The “AM CVn Trial” which used a mock courtroom procedure to debate the relative scientific merits of different kinds of WET targets, proved to be very entertaining but not very enlightening. If anything, it showed why the adversarial system of justice, at least as practiced in the United States, is not a very good way to get at truth or to sway opinion. We would be ill-advised, I think, to adopt it for any part of our future meetings. It was worth a try, though.

- Although much work remains to be done, it now looks like all of the white dwarf stars that exhibit coherent variations are open to quantitative analysis by asteroseismology. Observational and analytical techniques differ, but I think we have now identified at least one fruitful procedure for each of the known strips of instability along the white dwarf cooling sequence. This is quite an achievement, in my opinion, for the seven years the WET has been in operation.

- The superb spectroscopic data presented by Klaus Werner, from which he can derive accurate effective temperatures, makes an important contribution to the determination of asteroseismological distances, particularly for the hottest targets, where the bolometric correction, essential to turn the absolute luminosity into a distance, is large.

- When we established the Council of the Wise (COW) at the workshop in 1993, we agreed that the members would stand for re-election at our next workshop in 1995. The consensus to continue the current COW members until the workshop in 1997 suggests that whatever dissatisfaction with the current COW membership may exist, it is not sufficient to insist on any changes now. Following this decision the COW met and chose targets for two WET runs in 1996. As an experiment, the first 1996 WET run will be embedded in a much longer run with a Delta Scuti target, and will be controlled from Vienna, with support from Iowa and Texas. When results are in hand, the COW will decide on the success of the experiment.

- The timing errors uncovered during the re-reduction of the GD 358 WET data from 1990 are embarrassing, even though
they caused no real problems in the analysis, and we must work to eliminate them in the future. The most promising way to do this was demonstrated at the meeting by Denis Sullivan, who brought along a Global Positioning System (GPS) receiver and antenna.

• It is customary in a meeting summary to declare that the meeting was a success. I so declare.