Research Article

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A preliminary comparison of photometric (MWSC) and trigonometric (TGAS) distances of open cluster stars

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Abstract: The goal of this research was to compare the open cluster photometric distance scale of the global survey of star clusters in the Milky Way (MWSC) with the distances derived from trigonometric parallaxes from the Gaia DR1/TGAS catalogue and to investigate to which degree and extent both scales agree. We compared the parallax-based and photometry-based distances of 5743 cluster stars selected as members of 1118 clusters based on their kinematic and photometric MWSC membership probabilities. We found good overall agreement between trigonometric and photometric distances of open cluster stars. The residuals between them were small and unbiased up to $\log(d, [pc]) \approx 2.8$. If we considered only the most populated clusters and used cluster distances obtained from the mean trigonometric parallax of their MWSC members, the good agreement of the distance scales continued up to $\log(d, [pc]) \approx 3.3$.

Keywords: distance scale – parallaxes, distance scale - clusters

1 Introduction

The global survey Milky Way Star Clusters (MWSC; Kharchenko et al. 2012) is a comprehensive study of Galactic star clusters based on a combination of all-sky catalogues 2MASS (Skrutskie et al. 2006) and PPMXL (Röser et al. 2010). Within this project, the kinematic-photometric membership probabilities of stars in each cluster region and basic cluster parameters (including solar-centric distances) were determined for 2859 known (Kharchenko et al. 2013) and 202 newly-discovered (Schmeja et al. 2014; Scholz et al. 2015) open clusters. These clusters are widespread over the Galaxy (with solarcentric distances up to 15 kpc and the mode of 2.4 kpc) and are of great interest for structure and evolution studies of the Galactic disk population. Previous parallax-based cluster distances (Robichon et al. 1999; Loktin & Beshenov 2001) derived from Hipparcos parallaxes (ESA 1997) were given only for local objects and cannot be used for evaluating MWSC distances due to their insufficient extent and/or accuracy. The distances of MWSC clusters are based on the isochrone fitting technique and only indirectly (via evolutionary stellar models) related to the present cosmic distance scale.

The Gaia mission (Gaia Collaboration 2016), the Hipparcos successor, has all chances to improve in a few years, the situation by providing high precision data on trigonometric parallaxes of Galactic stars at distances up to 3 kpc from the Sun. Already with the Tycho-Gaia Astrometric Solution (TGAS) included in the first data release, Gaia DR1 (Michalik et al. 2015; Gaia Collaboration 2016, 2017; Arenou et al. 2017), parallax accuracies of the order of 0.3-0.7 mas were achieved for about 2 mln. stars. These allow us to compare the photometric distances of cluster stars directly with independent trigonometric measurements at distances up to the typical distance of MWSC clusters.

With this preliminary study (final results are published by Kovaleva et al. 2017), we provide a first comparison of the open cluster distance scale using photometric cluster distances from the MWSC survey and individual MWSC member distances derived from their TGAS trigonometric parallaxes. In Sect. 2, we describe the data used in this study. In Sect. 3, we perform the comparison and dis-
cuss the derived results of the study, and in Sect. 4, we draw basic conclusions following from the comparison.

2 The data

The input data for our comparison of distances to open cluster stars came from two independent sources: the MWSC and Gaia TGAS catalogues. The most probable members of MWSC star clusters were cross-identified with the TGAS catalogue. The mean photometric MWSC distances, derived from cluster Colour-Magnitude Diagrams (CMDs), were assigned to each member of a given cluster and then compared with their individual trigonometric distances derived from their TGAS parallaxes. Hereafter, we call the MWSC distances \( d_{MWSC} \), derived as a rule from CMD fitting, photometric distances, and those obtained from TGAS parallaxes \( d_{TGAS} \) trigonometric distances.

2.1 The MWSC: data description, limitations, distance scale, membership

The MWSC survey provides a comprehensive sample of star clusters of our Galaxy with a variety of well-determined parameters based on accurate photometric and uniform kinematic stellar data gathered from the all-sky catalogues 2MASS and PPMXL. The full sample contains 3208 objects: 3061 open and 147 globular clusters. In this study, we concentrated on Galactic disk objects (open clusters and associations) and refer to them as MWSC clusters. Within 2 kpc, corresponding to the distance one can reach with TGAS data, there were about 1500 MWSC clusters suitable for this study.

For our analysis, we selected stars with membership probabilities \( P > 60\% \). The probabilities were extracted from the MWSC catalogue of stars in cluster areas (Kharchenko et al. 2013; Schmeja et al. 2014; Scholz et al. 2015), where they were calculated for every star in the vicinity of the centre of a surveyed cluster based on photometric and kinematic criteria as described in (Kharchenko et al. 2012). They were used for the determination of basic cluster parameters such as the apparent size, mean proper motion, and distance. We refer to these selected stars as probable members, whereas we call those with even higher membership probabilities \( P > 90\% \) definite cluster members. The respective numbers of used cluster members are given in Table 1.

For all cluster members, we assigned the respective photometric distance of the cluster they belong to. Note that the expected variations of individual stellar distances within a cluster are relatively small. If one accepts that the typical cluster size is equal to the average tidal radius of MWSC clusters of about 10 pc (see Kharchenko et al. 2013), then the error would be about 10% at the smallest available MWSC distances of the order of 100 pc and less than 0.5% at about 2 kpc, corresponding to the average distance of MWSC clusters. This error lies well within the expected errors of TGAS parallaxes of 0.3-0.7 mas (Michalik et al. 2015), which is of the order of 100% at a distance of 2 kpc. The comparison of MWSC data with independent cluster distances from the literature (Kharchenko et al. 2013) yielded a typical error of MWSC distances of 11%.

2.2 Gaia TGAS: cross-identification, dataset characteristics

Gaia DR1 is the most precise all-sky astrometric survey since Hipparcos, providing TGAS parallaxes for 20 times more stars than known before and with unprecedented precision. However, according to Arenou et al. (2017), a global negative parallax zero point (about \(-0.04\) mas) is consistently found with independent estimation methods. These authors recommend to treat Gaia DR1 astrometric data accurately, taking into account a systematic error of about 0.3 mas in parallaxes. In spite of these limitations, Gaia DR1 provides a new basis for studies of Galactic structure, until it will be superseded by the coming Gaia DR2.

Our cross-identification of MWSC cluster members with TGAS stars was based on the Sky Algorithm of TopCat (Taylor 2005). In total, 5743 probable members of 1118 MWSC clusters were found to have TGAS parallaxes. The first ten of the most populated clusters contain more than 50 probable members with TGAS parallaxes each. For definite members, the numbers are 1058 stars and 481 clusters, respectively. Among them, negative TGAS parallaxes (\( \sigma_{TGAS} < 0 \)) were found for 199 probable, and 32 definite cluster members, respectively.

Figure 1 shows the distributions of relative parallax errors for probable and definite cluster members. Both distributions behave similarly, showing maxima at about \( \log \varepsilon_{\varpi}/\varpi = -0.5 \) and positive tails corresponding to up to two orders of magnitude larger errors, whereas on the other side the histogram extends to only one order of magnitude smaller errors. The similarity of the distributions implies that both samples are similar to each other with respect to their parallax quality. Therefore, we discuss mainly the probable cluster members.
smaller relative errors of the trigonometric parallaxes the distance range naturally shortens. However, the fraction of higher quality parallax stars off the agreement at large distances is larger. This speaks against the hypothesis of a Lutz-Kelker bias as the main reason for the bump, because this statistical effect critically depends on the relative parallax error becoming smaller with better quality parallaxes (see, for the case of a flat distribution of stars, Rastorguev et al. 2017). We considered these stars as foreground interlopers with correct \( \Delta d_{\text{TGAS}} \), but wrong membership probabilities (and hence their photometric distances). However, we noted that the fraction of such interlopers was small. We did also not see a significant difference in the scatter of the different groups of stars in Figure 2, when we used the definite cluster members (not shown) instead of the probable cluster members.

To estimate the degree of agreement between the photometric and trigonometric distances, one can use the relation between \( \Delta d = d_{\text{TGAS}} - d_{\text{MWSC}} \) and \( \log d_{\text{MWSC}} \) (Figure 3). The plot clearly confirms the impression that we already had from Figure 2. One can see that for \( \log d_{\text{MWSC}} \lesssim 2.8 \) the residuals are small and unbiased indicating very nice agreement of the two scales. At \( \Delta d < -1 \) kpc there is a long tail of negative residuals \( \Delta d \) , discussed above as possible foreground interlopers. These stars could be co-moving relatively bright foreground red giants or just affected by uncertain proper motions leading to wrong kinematic MWSC membership probabilities. In contrast, the
low-quality parallax dots clearly reside above zero indicating overestimated trigonometric distances due to a low-accuracy parallax bias.

Obviously, it should be also instructive to estimate the agreement between trigonometric and photometric distances for the clusters instead of their individual member stars. According to Arenou et al. (2017) and Arenou & Luri (1999), one should not average distances obtained from inverting observed parallaxes, but instead first average the parallaxes and then invert the result. Therefore, we first determined average TGAS parallaxes and their standard deviations for the members of every populated cluster. Then, we cleaned the datasets from values out of 3 standard deviations (suggesting them to be caused either by foreground co-moving stars or by low quality parallaxes, as discussed above). Our final TGAS cluster distances correspond to the average parallaxes of these cleaned datasets. Figure 4 represents the comparison of these distances with their photometric counterparts for the 65 most populated clusters containing 17 and more probable members with TGAS parallaxes, after the removal of outliers. The agreement of trigonometric and photometric cluster distances is much better than that for individual members and extends much further, up to 1.5...2 kpc, if we take into account only populated clusters. The bias is small, and the relation can be approximated by the weighted linear regression:

$$\log(d_{\text{TGAS}}) = a + b \log d_{\text{MWSC}},$$

where $a = -0.0245 \pm 0.0003$, $b = 1.0102 \pm 0.0001$.

4 Summary and conclusions

We reported the preliminary procedure and results of a direct comparison of the photometric distance scale derived within the MWSC survey with trigonometric distances available for probable members of MWSC clusters in the Gaia DR1 TGAS catalogue. In total, 5743 probable members of 1118 clusters were identified in the TGAS, with 1058 of them having the highest membership probabilities and definitely belonging to 489 clusters, according to the MWSC. We found that the trigonometric and photometric distances of cluster stars perfectly agree up to about 700 pc from the Sun. Beyond this limit, we observed a slowly increasing number of deviating trigonometric or photometric distances. Nevertheless, a good general agreement in both distance scales was seen up to 1.5...2 kpc from the Sun, if we used distances of whole clusters with a large number of member stars. Since there were less overestimations of the trigonometric distances, if we considered only the reliable parallaxes, we think that some differences in the distance scales are caused by the uncer-
tainty of trigonometric parallaxes. In case of clearly overestimated photometric distances, we attributed the deviation to wrong MWSC membership probabilities, mainly due to uncertain proper motions.

Note that the results of a further improved and final comparison of the MWSC and TGAS distance scales are published elsewhere (Kovaleva et al. 2017). In their analysis of the same input data, Kovaleva et al. (2017) make the comparison between TGAS and MWSC using parallaxes rather than distances. Their work involves photometric parallaxes defined as $\varpi_{\text{MWSC}} = 1000/d_{\text{MWSC}}$, where $\varpi_{\text{MWSC}}$ is in mas. The advantage is that these can be directly compared with all TGAS parallaxes, including negative values, to avoid any bias. The improved and extended study by Kovaleva et al. (2017) includes (1) an investigation of TGAS proper motions of MWSC cluster members, also used in the exclusion of obvious non-members, (2) a comparison of the MWSC and TGAS parallaxes of open clusters with the independent results from the Gaia investigation (Gaia Collaboration 2017) of nearby open clusters, (3) a closer look at some clusters with problematic MWSC parallaxes, and (4) an investigation of possible small systematic effects in the TGAS-MWSC parallax differences of open clusters as a function of Galactic longitude.

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References


